



## Article

# Intra-Night Variability of OJ 287 with Long-Term Multiband Optical Monitoring

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**Abstract:** We present long-term optical multi-band photometric monitoring of the blazar OJ 287 from 6 March 2010 to 3 April 2016, with high temporal resolution in the *VRI*-bands. The flux variations and colour-magnitude variations on long and short timescales were investigated to understand the emission mechanisms. In our observation, the major outbursts occurred in January 2016, as predicted by the binary pair of black holes model for OJ 287, with  $F_{var}$  of 1.3~2.1%, and variability amplitude (Amp) of 5.8~9.0%. The intra-night variability (IDV) durations were from 18.5 to 51.3 min, and the minimal variability timescale was about 4.7 min. The colour-magnitude variation showed a weak positive correlation on the long timescale with Pearson's  $r = 0.450$ , while a negative correlation was found on intra-night timescales. We briefly discuss the possible physical mechanisms that are most likely to be responsible for the observed flux and colour-magnitude correlation variability.

**Keywords:** BL Lacertae objects: general; quasar: individual: OJ 287; galaxies: photometry

## 1. Introduction

BL Lacertae (BL Lac) objects, which have either very weak or no emission lines [1], and flat-spectrum radio quasars (FSRQs) with strong emission lines [2,3] form a subclass of radio-loud active galactic nuclei (AGNs) known as blazars. Blazars are characterised by non-thermal emission, and strong and rapid flux variability across the entire electromagnetic spectrum, from radio- to gamma-rays. The emission is normally attributed to the relativistic jet oriented at a small angle from the line of sight [4]. Blazar flux variability timescales extending from a few minutes to years and even decades can be broadly divided into three classes: a large variation over hours to days is often known as intra-day or -night variability (IDV) or micro-variability [5,6]; variation on timescales of days to weeks, or even a few months, is considered short-term variability (STV); meanwhile, long-term variability (LTV) can have timescales from several months to years [7–10].

OJ 287 ( $\alpha = 08:54:48.9$ ,  $\delta = +20:06:30.6$ , J2000) is a blazar at redshift  $z = 0.306$  [11]. Sillanpää et al. [12] pointed out for the first time that there is a double-peak structure in the cyclic optical outbursts of OJ 287 by using the optical V-band observations starting from 1890. The curve exhibited periodic outbursts at intervals of  $\sim 12$  years. A binary pair of super-massive black holes system was used to explain this quasi-periodic light curve. Sillanpää et al. [13] (1996a [13], 1996b [14]) reported the next double-peak outburst occurrence in 1994–1995, which occurred almost exactly at the predicted times. Then a double-peaked outburst was also seen during the next recurrence in

2005–2008 [15]. These predicted recurrences are usually interpreted as OJ 287 housing a binary pair of black holes system with a period of  $\sim 12$  years.

In the case of OJ 287, outbursts occurring roughly every 12 years are almost certainly produced, and it was predicted that OJ 287 should show a major outburst in 2015–2016 [16]. In our work, along with the confirmation of this predicted outburst during 2015–2016 reported by [10,17], we study the properties of variability and spectral variation in a short timescales.

## 2. Observations and Data Reduction

Our photometric observations of blazar OJ 287 were performed using the 1.02 m (YAO1.02) optical telescopes at the Yunnan Astronomical Observatory. Since 2009, this telescope has been equipped with an Andor DW436 CCD camera ( $2048 \times 2048$  pixels) at the Cassegrain focus ( $f = 13.3$  m). The readout noise and gain are 6.33 electrons and 2.0 electrons/ADU in a  $2 \mu\text{s}$  readout time model, or 2.29 electrons and 1.4 electrons/ADU in a  $16 \mu\text{s}$  readout time model. The FOV of the CCD frame is  $7.3 \times 7.3 \text{ arcmin}^2$  with a pixel scale of  $0.21 \text{ arcsec/pixel}$ . We used standard Johnson broad-band filters with effective wavelength midpoints of  $V = 551 \text{ nm}$ ,  $R = 658 \text{ nm}$ , and  $I = 806 \text{ nm}$ .

We performed optical observations in the  $V$ -,  $R$ -, and  $I$ - bands in the cyclic mode. The exposure times from 50 to 300 s were chosen according to seeing, the filter and the telescope. The time resolutions (the time interval between two adjacent data points in the same band) were less than 15 min and were approximately 10 min in most cases. Therefore, these data in the cyclic mode were considered as quasi-simultaneous measurements, which were explored for analysing inter-band correlation and colour index. During our observation campaign (6 March 2010 to 3 April 2016), we observed for a total of 34 nights, to obtain 2255 CCD frames (shown in Table 1).

**Table 1.** A short summary of the observations in each band performed from 2010 and 2016. In total we observed for 34 nights and obtained 2255 CCD frames.

Year	Nights	V(N)	R(N)	I(N)
2010	10	183	196	197
2012	4	133	136	136
2013	7	142	153	156
2014	5	42	46	55
2015	2	19	25	39
2016	6	199	201	197

All images were reduced with standard image reduction and analysis facility (IRAF)<sup>1</sup> procedures after bias and flat-field corrections. Aperture photometry was performed on the source and comparison stars with APPHOT. Photometry of the source and comparison stars was performed with the same aperture, which was determined by the full width at half-maximum (FWHM) of the comparison stars and was the same for each observation. We compared the photometric results and found that an aperture radius of about 1.5 FWHM almost always provided the best S/N ratio, and thus we concentrated on this aperture for our analysis. Each night, we performed aperture photometry with different aperture radii.

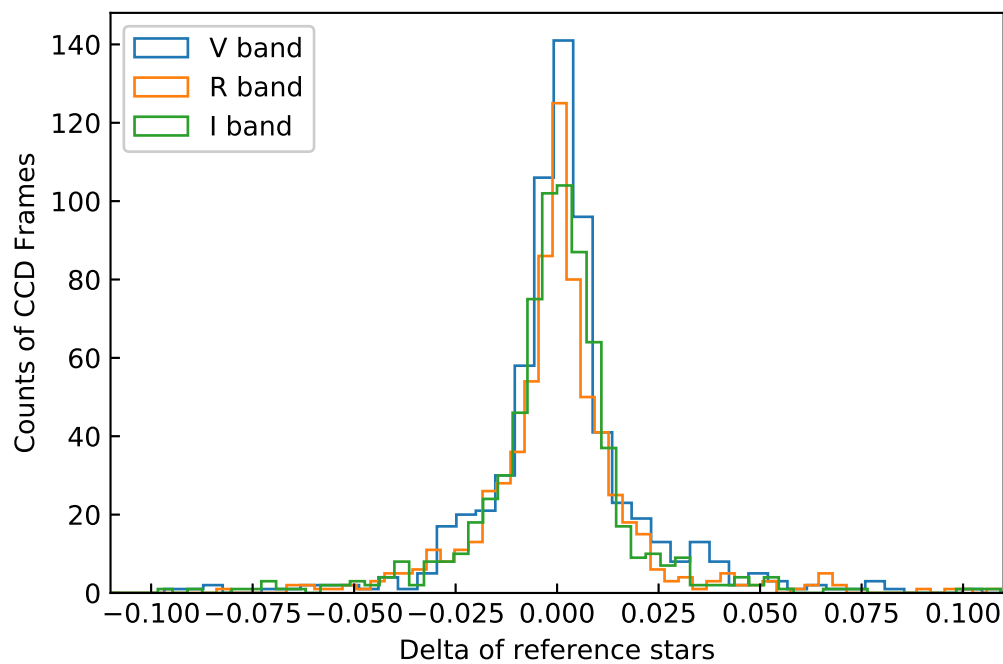
For each CCD frame, the instrumental magnitudes of OJ 287 and the two comparison stars (listed in Table 2) were extracted. The brightness (magnitude) of OJ 287 was calculated as the average of that derived with respect to comparison stars 4 and 10, and the corresponding standard deviation was treated as the error of OJ 287 within each CCD frame. The deviation of the average of the differential instrumental magnitude of comparison stars 4 and 10,  $\Delta(\text{star 4} - \text{star 10})$ , was used to verify the

<sup>1</sup> IRAF is distributed by the National Optical Astronomy Observatories, which are operated by the Association of Universities for Research in Astronomy, Inc., under cooperative agreement with the National Science Foundation.

stability of the comparison stars, which is also taken as the accuracy of the observations. Figure 1 shows the deviation of comparison stars 4 and 10 for the entire observational campaign. The deviations in bands *V*-, *R*- and *I*-bands are mainly distributed in  $\pm 0.025$  mag.

**Table 2.** The magnitudes of the comparison stars. Comparison stars (VRI) from Fiorucci & Tosti [18].

Star ID	V (mag)	R (mag)	I (mag)
4	14.18 (0.04)	13.74 (0.04)	13.28 (0.04)
10	14.60 (0.05)	14.34 (0.05)	14.03 (0.05)



**Figure 1.** Distribution of the deviation of the average of the differential instrumental magnitudes of comparison stars 4 and 10 in bands *V*, *R* and *I*.

### 3. Results

#### 3.1. Variability

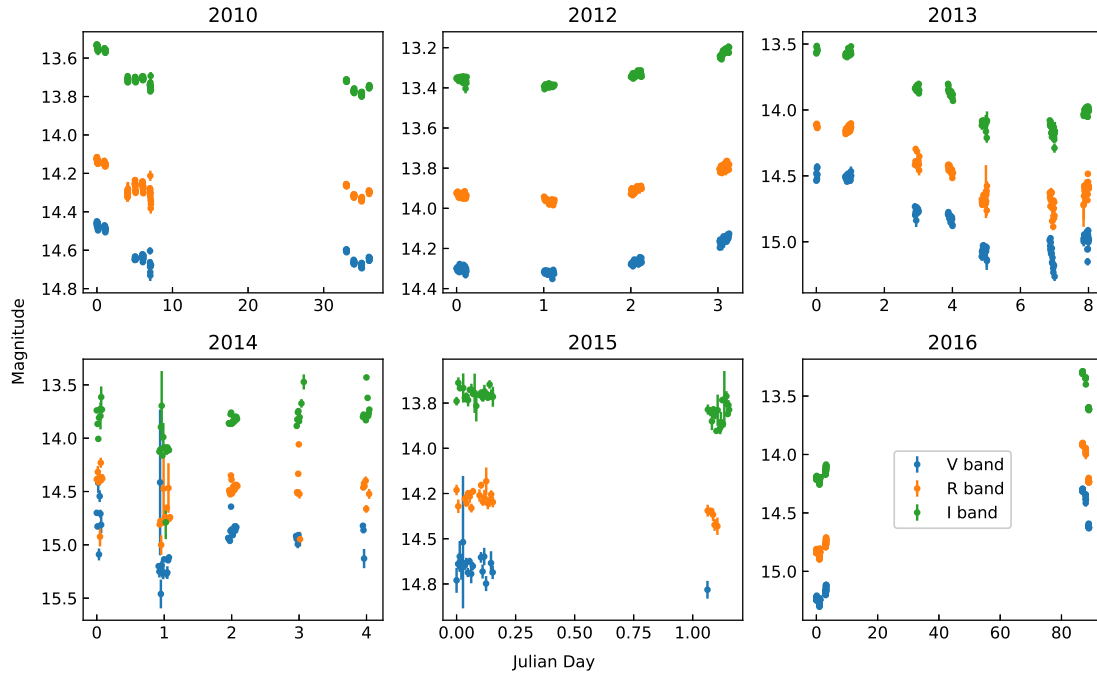
The *V*, *R* and *I* light curves from 6 March 2010 to 3 April 2016 are shown in Figure 2. The average magnitudes in each band are  $\bar{V} = 14.728 \pm 0.352$ ,  $\bar{R} = 14.351 \pm 0.328$  and  $\bar{I} = 13.763 \pm 0.309$ . The amplitude variability (maximum minus minimum) in each filter are respectively  $\Delta V = 1.^m335$ ,  $\Delta R = 1.^m237$  and  $\Delta I = 1.^m594$ .

Presently, a number of statistical tests, such as the *C*-test, the *F*-test, the  $\chi^2$ -test and one-way analysis of variance (ANOVA), have been proposed to assess quantitatively whether there are IDVs [6,9,19–23]. Here, we apply the *F*-test and  $\chi^2$ -test to cross-check the intra-day light curves. We consider variability when the light curve satisfied both criteria described below.

The *F*-test is regarded as a powerful distribution statistic to check for the presence of variability, as introduced by [21]. When comparing two sample variances, the *F*-statistic values are calculated as:

$$F_1 = \frac{S_{BL-StarA}^2}{S_{StarA-StarB}^2}, F_2 = \frac{S_{BL-StarB}^2}{S_{StarA-StarB}^2} \quad (1)$$

where  $S_{BL-StarA}^2$ ,  $S_{BL-StarB}^2$  and  $S_{StarA-StarB}^2$  are the variances of the differential instrumental magnitudes of the blazar and comparison star A, the blazar and comparison star B, and comparison stars A and B, respectively. The  $F$ -statistic is compared with the  $F_{\nu_{BL}, \nu_c}^{(\alpha)}$  critical value. The number of degrees of freedom for each sample,  $\nu_{BL}$  and  $\nu_c$ , are the number of measurements minus 1 ( $N - 1$ ), and  $\alpha$  is the significance level set for the test. In this paper, the  $F$ -test was performed at two significance levels: 0.99 and 0.999. If either  $F_1$  or  $F_2$  was larger than the critical value, the null hypothesis (no variability) was discarded.

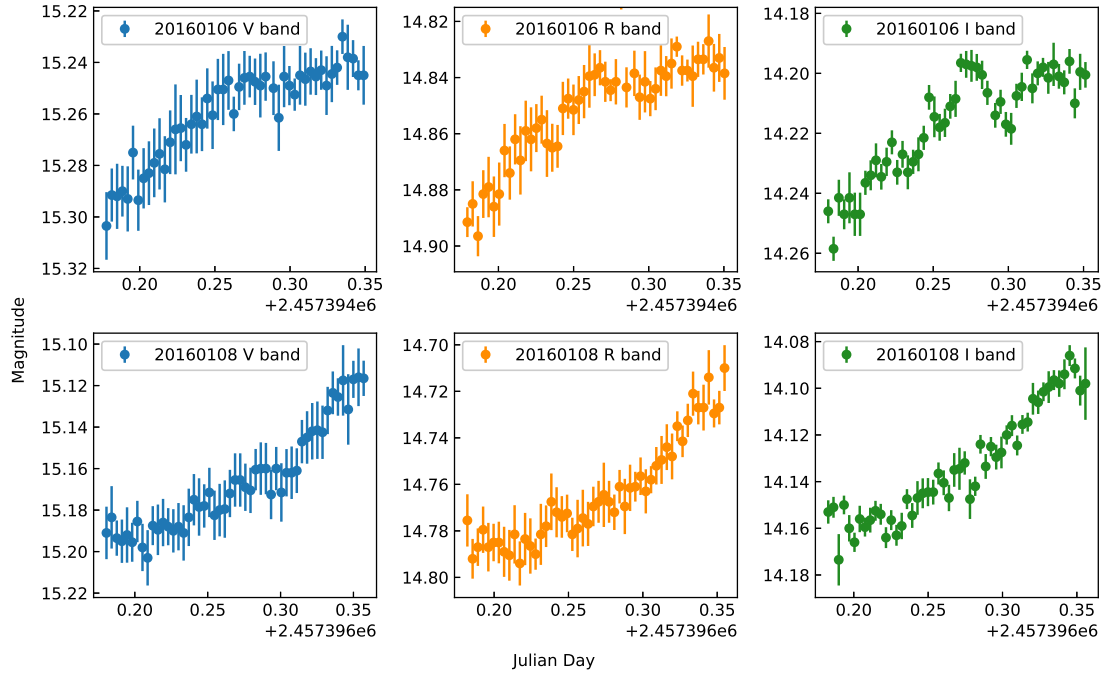


**Figure 2.** Light curves from 6 March 2010 to 3 April 2016 in V-, R- and I-bands. The error bars of some data points are smaller than the symbol.

The  $\chi^2$ -test was also used:

$$\chi^2 = \sum_{i=1}^N \left( \frac{m_i - \bar{m}}{\sigma_{m_i}} \right)^2 \quad (2)$$

where  $m_i$  are the individual magnitudes,  $\sigma_{m_i}$  are their errors, and  $\bar{m}$  denotes the mean value of magnitude for an observation night. This statistic is also compared against the critical value  $\chi_{\alpha, \nu}^2$  obtained from the  $\chi^2$  probability function, where  $\nu$  is the degrees of freedom and  $\alpha$  is the significance level.  $\chi^2 > \chi_{\alpha, \nu}^2$  provides evidence of variability. OJ 287 was said to be variable (V) if the statistic from both tests was satisfied at the 0.999 level, while it was marked as “non-variable” (NV) if none of these criteria were met at the 0.99 level. For neither, it was marked as “probably variable” (PV). The results of the  $F$ - and  $\chi^2$ -test and the IDV observations in the V-, R- and I-bands are presented in Table 3. As shown in Table 3, 17 light curves were marked as variable among 100 light curves (18 light curves met the criterion with the  $F$ -test and 57 light curves with the  $\chi^2$ -test). As an illustration, the significant variability on two nights are presented, 6–8 January in 2016, which were marked as variable in the V-, R- and I-bands (shown in Figure 3). On January 6, 2016, the magnitudes showed a monotonically increasing trend. The magnitude changes of the V-, R- and I-bands were  $\Delta V = 0.^m074$ ,  $\Delta R = 0.^m072$  and  $\Delta I = 0.^m063$  in 252 min. On January 8, 2016, the same situation occurred; magnitudes continued to increase by  $\Delta V = 0.^m087$ ,  $\Delta R = 0.^m084$  and  $\Delta I = 0.^m088$  in 254 min.



**Figure 3.** Light curves of two significant micro-variabilities on 6–8 January in 2016, which were marked as variable in V-, R- and I-bands.

For each IDV, the variability amplitude (Amp) could be calculated by [24]:

$$Amp = \sqrt{(A_{max} - A_{min})^2 - 2\sigma^2} \quad (3)$$

where  $A_{max}$  and  $A_{min}$  are the maximum and minimum magnitude, respectively, of the light curve for the night being considered, and  $\sigma$  is the corresponding standard deviation.

Our observational data were obtained with different exposures and time intervals across 6 years. To avoid uncertainties caused by different observational set-ups, the fractional variability amplitude ( $F_{var}$ ) was calculated to show the intrinsic variability amplitude of the source by removing the effects of measurement noise; it is defined as follows [25]:

$$F_{var} = \frac{1}{\bar{m}} \sqrt{S^2 - \overline{\sigma_{err}^2}} \quad (4)$$

Here  $S^2$  denotes the total variance of the light curve:

$$S^2 = \frac{1}{N-1} \sum_{i=1}^N (m_i - \bar{m})^2 \quad (5)$$

where  $\overline{\sigma_{err}^2}$  is the mean error squared, and  $\bar{m}$  is the mean value of the magnitude. The error on  $F_{var}$  is

$$\sigma_{F_{var}} = \frac{1}{2F_{var}} \sqrt{\frac{1}{N} \frac{S^2}{\bar{m}^2}} \quad (6)$$

**Table 3.** Results of the intra-day/night variability (IDV) observations of OJ 287.

Date	Band	N	Magnitude (mag)	F-Test	$\chi^2$ -Test	Amp %	Flux mJy	$F_{var}$ %	State
				$F_1, F_2, F_{0.99}, F_{0.999}$	$\chi^2, \chi^2_{0.99}, \chi^2_{0.999}$				
6 March 2010	V	33	14.468 $\pm$ 0.011	2.996, 1.858, 2.318, 3.092	219.915, 53.486, 62.487	4.480	6.172 $\pm$ 0.063	0.908 $\pm$ 0.001	PV
	R	33	14.136 $\pm$ 0.009	1.790, 2.214, 2.318, 3.092	42.286, 53.486, 62.487	3.205	6.520 $\pm$ 0.054	0.000 $\pm$ 0.000	PV
	I	33	13.543 $\pm$ 0.010	2.039, 0.685, 2.318, 3.092	79.848, 53.486, 62.487	3.284	10.079 $\pm$ 0.088	0.444 $\pm$ 0.002	PV
7 March 2010	V	25	14.485 $\pm$ 0.009	1.783, 1.285, 2.659, 3.735	144.765, 42.980, 51.179	3.372	6.075 $\pm$ 0.050	0.653 $\pm$ 0.001	PV
	R	25	14.152 $\pm$ 0.007	0.798, 1.501, 2.659, 3.735	21.153, 42.980, 51.179	2.748	6.425 $\pm$ 0.039	0.000 $\pm$ 0.000	NV
	I	25	13.559 $\pm$ 0.005	0.709, 0.365, 2.659, 3.735	34.276, 42.980, 51.179	2.273	9.932 $\pm$ 0.050	0.000 $\pm$ 0.000	NV
10 March 2010	R	12	14.301 $\pm$ 0.014	0.936, 0.502, 4.462, 7.761	3.314, 24.725, 31.264	4.001	5.600 $\pm$ 0.071	0.000 $\pm$ 0.000	NV
	I	12	13.713 $\pm$ 0.009	1.427, 0.706, 4.462, 7.761	46.593, 24.725, 31.264	2.900	8.625 $\pm$ 0.069	0.562 $\pm$ 0.002	PV
11 March 2010	V	20	14.643 $\pm$ 0.005	1.146, 0.579, 3.027, 4.474	23.267, 36.191, 43.820	1.568	5.254 $\pm$ 0.027	0.000 $\pm$ 0.000	NV
	R	19	14.268 $\pm$ 0.020	7.582, 8.658, 3.128, 4.683	309.459, 34.805, 42.312	6.029	5.773 $\pm$ 0.105	1.740 $\pm$ 0.003	V
	I	20	13.712 $\pm$ 0.007	1.382, 1.689, 3.027, 4.474	17.720, 36.191, 43.820	2.597	8.634 $\pm$ 0.059	0.000 $\pm$ 0.000	NV
12 March 2010	V	26	14.636 $\pm$ 0.009	2.454, 2.456, 2.604, 3.629	33.089, 44.314, 52.620	3.879	5.287 $\pm$ 0.046	0.376 $\pm$ 0.003	NV
	R	26	14.263 $\pm$ 0.016	1.599, 3.275, 2.604, 3.629	128.718, 44.314, 52.620	5.727	5.799 $\pm$ 0.084	1.275 $\pm$ 0.002	PV
	I	26	13.704 $\pm$ 0.006	1.287, 1.618, 2.604, 3.629	39.641, 44.314, 52.620	1.709	8.690 $\pm$ 0.047	0.265 $\pm$ 0.002	NV
13 March 2010	V	11	14.679 $\pm$ 0.030	1.028, 2.299, 4.849, 8.754	29.427, 23.209, 29.588	12.060	5.083 $\pm$ 0.144	2.381 $\pm$ 0.007	PV
	R	13	14.308 $\pm$ 0.042	3.205, 6.351, 4.155, 7.005	121.479, 26.217, 32.909	15.897	5.566 $\pm$ 0.215	3.605 $\pm$ 0.008	PV
	I	13	13.744 $\pm$ 0.019	1.566, 3.960, 4.155, 7.005	63.668, 26.217, 32.909	7.597	8.382 $\pm$ 0.147	1.523 $\pm$ 0.004	PV
8 April 2010	V	16	14.602 $\pm$ 0.004	0.781, 1.767, 3.522, 5.535	2.780, 30.578, 37.697	1.395	5.454 $\pm$ 0.020	0.000 $\pm$ 0.000	NV
	R	16	14.261 $\pm$ 0.003	1.034, 0.769, 3.522, 5.535	3.770, 30.578, 37.697	1.015	5.807 $\pm$ 0.016	0.000 $\pm$ 0.000	NV
	I	16	13.717 $\pm$ 0.004	0.662, 1.544, 3.522, 5.535	30.781, 30.578, 37.697	1.199	8.587 $\pm$ 0.028	0.197 $\pm$ 0.001	PV
9 April 2010	V	16	14.664 $\pm$ 0.007	1.056, 1.612, 3.522, 5.535	17.115, 30.578, 37.697	2.216	5.154 $\pm$ 0.031	0.000 $\pm$ 0.000	NV
	R	16	14.317 $\pm$ 0.005	1.194, 1.316, 3.522, 5.535	11.264, 30.578, 37.697	1.662	5.517 $\pm$ 0.025	0.000 $\pm$ 0.000	NV
	I	16	13.772 $\pm$ 0.006	1.201, 3.118, 3.522, 5.535	39.401, 30.578, 37.697	2.344	8.165 $\pm$ 0.046	0.416 $\pm$ 0.001	PV
10 April 2010	V	21	14.680 $\pm$ 0.008	0.991, 1.034, 2.938, 4.290	31.456, 37.566, 45.315	2.504	5.079 $\pm$ 0.038	0.000 $\pm$ 0.000	NV
	R	21	14.333 $\pm$ 0.006	0.987, 1.140, 2.938, 4.290	19.336, 37.566, 45.315	2.162	5.434 $\pm$ 0.028	0.000 $\pm$ 0.000	NV
	I	21	13.786 $\pm$ 0.007	1.137, 1.657, 2.938, 4.290	30.316, 37.566, 45.315	2.185	8.060 $\pm$ 0.052	0.000 $\pm$ 0.000	NV

Table 3. Cont.

Date	Band	N	Magnitude (mag)	F-Test	$\chi^2$ -Test	Amp %	Flux mJy	$F_{var}$ %	State
				$F_1, F_2, F_{0.99}, F_{0.999}$	$\chi^2, \chi^2_{0.99}, \chi^2_{0.999}$				
11 April 2010	V	15	$14.644 \pm 0.007$	0.816, 1.661, 3.698, 5.930	13.450, 29.141, 36.123	1.776	$5.247 \pm 0.031$	$0.000 \pm 0.000$	NV
	R	15	$14.297 \pm 0.004$	0.923, 0.620, 3.698, 5.930	8.547, 29.141, 36.123	1.651	$5.620 \pm 0.021$	$0.000 \pm 0.000$	NV
	I	15	$13.749 \pm 0.005$	1.502, 0.635, 3.698, 5.930	18.913, 29.141, 36.123	1.653	$8.341 \pm 0.039$	$0.113 \pm 0.004$	NV
12 April 2012	V	35	$14.301 \pm 0.011$	0.867, 1.752, 2.258, 2.983	46.245, 56.061, 65.247	5.208	$7.197 \pm 0.075$	$0.000 \pm 0.000$	NV
	R	36	$13.936 \pm 0.010$	0.781, 1.054, 2.231, 2.934	49.662, 57.342, 66.619	3.440	$7.834 \pm 0.070$	$0.000 \pm 0.000$	NV
	I	37	$13.358 \pm 0.012$	1.144, 0.759, 2.205, 2.888	64.849, 58.619, 67.985	5.876	$11.959 \pm 0.126$	$0.627 \pm 0.002$	PV
13 April 2012	V	24	$14.322 \pm 0.010$	0.830, 1.150, 2.719, 3.853	26.900, 41.638, 49.728	4.340	$7.060 \pm 0.063$	$0.000 \pm 0.000$	NV
	R	25	$13.964 \pm 0.010$	1.058, 2.296, 2.659, 3.735	38.305, 42.980, 51.179	4.025	$7.636 \pm 0.068$	$0.234 \pm 0.005$	PV
	I	25	$13.389 \pm 0.006$	0.494, 0.577, 2.659, 3.735	28.780, 42.980, 51.179	2.775	$11.616 \pm 0.064$	$0.000 \pm 0.000$	NV
14 April 2012	V	38	$14.268 \pm 0.010$	0.925, 0.927, 2.181, 2.844	35.374, 59.893, 69.346	4.220	$7.417 \pm 0.068$	$0.000 \pm 0.000$	PV
	R	39	$13.906 \pm 0.010$	1.039, 0.903, 2.157, 2.803	58.151, 61.162, 70.703	4.463	$8.054 \pm 0.077$	$0.000 \pm 0.000$	NV
	I	39	$13.333 \pm 0.010$	0.938, 1.674, 2.157, 2.803	96.397, 61.162, 70.703	3.846	$12.235 \pm 0.113$	$0.658 \pm 0.001$	PV
15 April 2012	V	36	$14.157 \pm 0.016$	0.763, 1.056, 2.231, 2.934	114.645, 57.342, 66.619	6.646	$8.219 \pm 0.117$	$0.650 \pm 0.004$	PV
	R	36	$13.795 \pm 0.016$	1.698, 1.443, 2.231, 2.934	123.812, 57.342, 66.619	5.563	$8.922 \pm 0.130$	$1.025 \pm 0.002$	PV
	I	35	$13.224 \pm 0.015$	1.085, 1.146, 2.258, 2.983	150.305, 56.061, 65.247	5.932	$13.528 \pm 0.186$	$1.048 \pm 0.002$	PV
16 March 2013	V	9	$14.484 \pm 0.038$	2.330, 0.912, 6.029, 12.046	29.022, 20.090, 26.124	9.137	$6.085 \pm 0.213$	$2.830 \pm 0.010$	PV
	R	9	$14.119 \pm 0.010$	0.116, 0.919, 6.029, 12.046	3.322, 20.090, 26.124	2.837	$6.620 \pm 0.064$	$0.000 \pm 0.000$	NV
	I	9	$13.546 \pm 0.018$	0.488, 1.091, 6.029, 12.046	27.032, 20.090, 26.124	5.120	$10.056 \pm 0.165$	$0.877 \pm 0.007$	PV
17 March 2013	V	27	$14.509 \pm 0.018$	0.584, 1.063, 2.554, 3.532	34.531, 45.642, 54.052	7.448	$5.942 \pm 0.096$	$0.000 \pm 0.000$	NV
	R	29	$14.149 \pm 0.019$	1.534, 2.097, 2.464, 3.362	63.400, 48.278, 56.892	7.599	$6.444 \pm 0.112$	$0.882 \pm 0.005$	PV
	I	29	$13.569 \pm 0.018$	0.474, 1.714, 2.464, 3.362	71.269, 48.278, 56.892	7.763	$9.843 \pm 0.160$	$1.176 \pm 0.003$	PV
19 March 2013	V	14	$14.771 \pm 0.027$	0.877, 1.384, 3.905, 6.409	15.011, 27.688, 34.528	9.850	$4.669 \pm 0.114$	$1.090 \pm 0.010$	NV
	R	17	$14.395 \pm 0.039$	2.594, 4.396, 3.372, 5.205	74.936, 32.000, 39.252	15.179	$5.138 \pm 0.189$	$3.168 \pm 0.007$	PV
	I	17	$13.837 \pm 0.017$	0.762, 0.926, 3.372, 5.205	31.938, 32.000, 39.252	6.710	$7.691 \pm 0.124$	$0.801 \pm 0.006$	NV
20 March 2013	V	22	$14.820 \pm 0.026$	2.317, 2.289, 2.857, 4.127	57.574, 38.932, 46.797	9.760	$4.464 \pm 0.108$	$1.803 \pm 0.005$	PV
	R	23	$14.454 \pm 0.024$	4.800, 3.911, 2.785, 3.983	57.654, 40.289, 48.268	9.423	$4.863 \pm 0.106$	$1.564 \pm 0.004$	V
	I	24	$13.871 \pm 0.028$	2.890, 3.196, 2.719, 3.853	189.788, 41.638, 49.728	12.291	$7.458 \pm 0.191$	$2.310 \pm 0.004$	PV

Table 3. Cont.

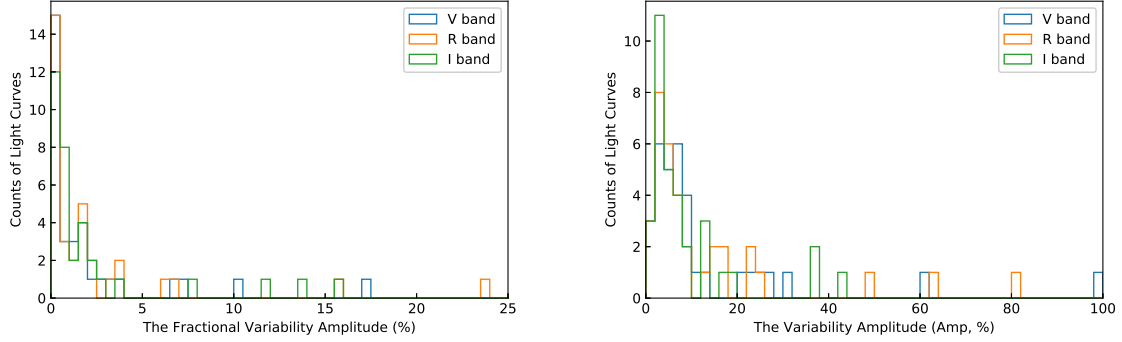
Date	Band	N	Magnitude (mag)	F-Test	$\chi^2$ -Test	Amp %	Flux mJy	$F_{var}$ %	State
				$F_1, F_2, F_{0.99}, F_{0.999}$	$\chi^2, \chi^2_{0.99}, \chi^2_{0.999}$				
21 March 2013	V	24	$15.074 \pm 0.031$	1.165, 0.624, 2.719, 3.853	38.513, 41.638, 49.728	11.054	$3.533 \pm 0.101$	$1.050 \pm 0.011$	NV
	R	26	$14.680 \pm 0.038$	1.070, 0.460, 2.604, 3.629	64.444, 44.314, 52.620	17.809	$3.951 \pm 0.140$	$0.000 \pm 0.000$	PV
	I	26	$14.103 \pm 0.028$	1.161, 0.612, 2.604, 3.629	32.312, 44.314, 52.620	12.661	$6.022 \pm 0.148$	$0.958 \pm 0.009$	NV
23 March 2013	V	18	$15.102 \pm 0.086$	4.072, 4.066, 3.242, 4.924	155.018, 33.409, 40.790	26.438	$3.451 \pm 0.271$	$7.236 \pm 0.014$	V
	R	21	$14.716 \pm 0.077$	4.759, 4.914, 2.938, 4.290	197.476, 37.566, 45.315	24.118	$3.831 \pm 0.265$	$6.347 \pm 0.012$	V
	I	23	$14.153 \pm 0.048$	1.671, 1.774, 2.785, 3.983	112.549, 40.289, 48.268	19.815	$5.755 \pm 0.250$	$3.583 \pm 0.008$	PV
24 March 2013	V	28	$14.981 \pm 0.045$	3.927, 3.201, 2.507, 3.443	97.643, 46.963, 55.476	22.681	$3.851 \pm 0.153$	$3.383 \pm 0.006$	V
	R	28	$14.598 \pm 0.048$	0.224, 1.097, 2.507, 3.443	116.079, 46.963, 55.476	22.819	$4.262 \pm 0.185$	$2.327 \pm 0.011$	PV
	I	28	$14.012 \pm 0.022$	1.513, 2.013, 2.507, 3.443	103.345, 46.963, 55.476	7.233	$6.547 \pm 0.136$	$1.623 \pm 0.004$	PV
31 March 2014	V	7	$14.729 \pm 0.199$	4.712, 8.024, 8.466, 20.030	81.483, 16.812, 22.458	60.883	$4.936 \pm 0.891$	$17.063 \pm 0.051$	PV
	R	7	$14.430 \pm 0.209$	56.965, 52.442, 8.466, 20.030	72.512, 16.812, 22.458	62.691	$5.053 \pm 0.824$	$15.795 \pm 0.045$	V
	I	8	$13.786 \pm 0.108$	1.572, 0.982, 6.993, 15.019	297.725, 18.475, 24.322	36.097	$8.102 \pm 0.785$	$7.760 \pm 0.030$	PV
1 April 2014	V	12	$15.164 \pm 0.242$	0.029, 1.201, 4.462, 7.761	34.159, 24.725, 31.264	98.843	$3.347 \pm 0.974$	$0.000 \pm 0.000$	PV
	R	11	$14.725 \pm 0.144$	0.091, 1.260, 4.849, 8.754	87.322, 23.209, 29.588	49.352	$3.824 \pm 0.521$	$2.436 \pm 0.163$	PV
	I	13	$14.105 \pm 0.231$	0.717, 2.427, 4.155, 7.005	166.964, 26.217, 32.909	104.231	$6.133 \pm 1.159$	$13.975 \pm 0.050$	PV
2 April 2014	V	15	$14.869 \pm 0.074$	11.715, 10.409, 3.698, 5.930	324.871, 29.141, 36.123	30.330	$4.277 \pm 0.310$	$6.913 \pm 0.014$	V
	R	15	$14.458 \pm 0.045$	10.458, 10.827, 3.698, 5.930	204.227, 29.141, 36.123	16.505	$4.849 \pm 0.205$	$3.999 \pm 0.008$	V
	I	15	$13.826 \pm 0.033$	3.458, 2.949, 3.698, 5.930	291.375, 29.141, 36.123	9.400	$7.772 \pm 0.238$	$2.869 \pm 0.006$	PV
3 April 2014	V	4	$14.941 \pm 0.033$	0.625, 1.733, 29.457, 141.108	5.923, 11.345, 16.266	7.140	$3.994 \pm 0.121$	$1.760 \pm 0.018$	NV
	R	6	$14.481 \pm 0.265$	59.894, 73.206, 10.967, 29.752	1001.118, 15.086, 20.515	80.609	$4.884 \pm 1.174$	$23.894 \pm 0.070$	V
	I	8	$13.750 \pm 0.121$	8.107, 3.941, 6.993, 15.019	443.834, 18.475, 24.322	37.464	$8.385 \pm 1.006$	$11.503 \pm 0.031$	PV
4 April 2014	V	4	$14.958 \pm 0.124$	6.919, 2.946, 29.457, 141.108	50.771, 11.345, 16.266	25.301	$3.954 \pm 0.443$	$10.089 \pm 0.044$	PV
	R	7	$14.478 \pm 0.084$	2.329, 4.181, 8.466, 20.030	49.708, 16.812, 22.458	23.629	$4.772 \pm 0.350$	$6.598 \pm 0.022$	PV
	I	11	$13.701 \pm 0.155$	101.854, 98.563, 4.849, 8.754	1163.937, 23.209, 29.588	42.137	$8.815 \pm 1.379$	$15.563 \pm 0.034$	V
27 March 2015	V	17	$14.650 \pm 0.053$	0.129, 0.562, 3.372, 5.205	18.000, 32.000, 39.252	21.636	$5.228 \pm 0.258$	$0.000 \pm 0.000$	NV
	R	17	$14.270 \pm 0.039$	0.431, 1.543, 3.372, 5.205	35.530, 32.000, 39.252	13.775	$5.766 \pm 0.210$	$1.923 \pm 0.012$	PV
	I	21	$13.699 \pm 0.030$	0.401, 0.751, 2.938, 4.290	23.917, 37.566, 45.315	12.073	$8.736 \pm 0.242$	$0.000 \pm 0.000$	NV



Table 3. Cont.

Date	Band	N	Magnitude (mag)	F-Test	$\chi^2$ -Test	Amp %	Flux mJy	$F_{var}$ %	State
				$F_1, F_2, F_{0.99}, F_{0.999}$	$\chi^2, \chi^2_{0.99}, \chi^2_{0.999}$				
28 March 2015	V	2	$14.715 \pm 0.068$	—	—	9.546	$4.928 \pm 0.306$	$3.685 \pm 0.052$	—
	R	8	$14.384 \pm 0.034$	0.439, 0.354, 6.993, 15.019	10.458, 18.475, 24.322	7.135	$5.187 \pm 0.162$	$0.000 \pm 0.000$	NV
	I	18	$13.802 \pm 0.053$	0.800, 0.458, 3.242, 4.924	95.727, 33.409, 40.790	17.694	$7.954 \pm 0.384$	$0.970 \pm 0.040$	PV
5 January 2016	V	49	$15.233 \pm 0.009$	3.775, 5.041, 1.977, 2.489	39.766, 73.683, 84.037	4.124	$3.052 \pm 0.024$	$0.000 \pm 0.000$	PV
	R	49	$14.830 \pm 0.009$	3.676, 3.784, 1.977, 2.489	104.805, 73.683, 84.037	3.368	$3.441 \pm 0.029$	$0.527 \pm 0.001$	V
	I	49	$14.201 \pm 0.007$	1.467, 1.474, 1.977, 2.489	109.280, 73.683, 84.037	3.264	$5.500 \pm 0.034$	$0.413 \pm 0.001$	PV
6 January 2016	V	50	$15.260 \pm 0.018$	11.933, 15.218, 1.963, 2.465	165.498, 74.919, 85.351	6.872	$2.975 \pm 0.050$	$1.346 \pm 0.002$	V
	R	50	$14.852 \pm 0.020$	17.948, 18.593, 1.963, 2.465	395.483, 74.919, 85.351	8.972	$3.371 \pm 0.061$	$1.635 \pm 0.002$	V
	I	50	$14.218 \pm 0.018$	8.024, 10.653, 1.963, 2.465	771.064, 74.919, 85.351	5.789	$5.417 \pm 0.087$	$1.547 \pm 0.002$	V
8 January 2016	V	52	$15.167 \pm 0.024$	37.909, 37.202, 1.936, 2.419	251.989, 77.386, 87.968	7.998	$3.243 \pm 0.073$	$1.956 \pm 0.003$	V
	R	51	$14.764 \pm 0.022$	25.587, 28.567, 1.949, 2.441	340.977, 76.154, 86.661	7.788	$3.655 \pm 0.076$	$1.870 \pm 0.002$	V
	I	51	$14.134 \pm 0.023$	9.004, 10.514, 1.949, 2.441	1089.284, 76.154, 86.661	8.108	$5.852 \pm 0.126$	$2.084 \pm 0.002$	V
1 April 2016	V	12	$14.305 \pm 0.007$	1.185, 0.758, 4.462, 7.761	7.032, 24.725, 31.264	2.439	$7.171 \pm 0.048$	$0.000 \pm 0.000$	NV
	R	12	$13.912 \pm 0.007$	1.625, 1.915, 4.462, 7.761	11.356, 24.725, 31.264	2.190	$8.010 \pm 0.051$	$0.000 \pm 0.000$	NV
	I	11	$13.297 \pm 0.007$	3.768, 4.057, 4.849, 8.754	231.306, 23.209, 29.588	2.148	$12.651 \pm 0.078$	$0.577 \pm 0.001$	PV
2 April 2016	V	10	$14.364 \pm 0.024$	1.266, 2.796, 5.351, 10.107	93.338, 21.666, 27.877	6.432	$6.793 \pm 0.151$	$1.516 \pm 0.007$	PV
	R	10	$13.969 \pm 0.021$	1.404, 0.485, 5.351, 10.107	192.535, 21.666, 27.877	4.774	$7.604 \pm 0.145$	$0.602 \pm 0.013$	PV
	I	7	$13.353 \pm 0.020$	0.863, 2.103, 8.466, 20.030	84.205, 16.812, 22.458	5.465	$12.017 \pm 0.217$	$1.561 \pm 0.006$	PV
3 April 2016	V	26	$14.616 \pm 0.009$	0.253, 1.588, 2.604, 3.629	32.455, 44.314, 52.620	3.759	$5.385 \pm 0.043$	$0.000 \pm 0.000$	NV
	R	29	$14.226 \pm 0.008$	0.412, 1.525, 2.464, 3.362	42.732, 48.278, 56.892	3.556	$5.997 \pm 0.047$	$0.246 \pm 0.003$	NV
	I	29	$13.609 \pm 0.006$	0.517, 0.920, 2.464, 3.362	80.325, 48.278, 56.892	2.370	$9.491 \pm 0.049$	$0.239 \pm 0.001$	PV

Our results of the IDV analysis are shown in Figure 4 and Table 3. Both the statistical distributions of two variability amplitudes show that the IDV amplitudes of OJ 287 during the whole observational campaign were small.



**Figure 4.** Distribution of the fractional variability amplitude ( $F_{var}$ ) and variability amplitude (Amp) in V, R and I in Table 3.

We extracted the variability for candidate light curves (marked as PV or V states) by an asymmetric flare template [26], which describes flares as exponentially rising and decaying by the function:

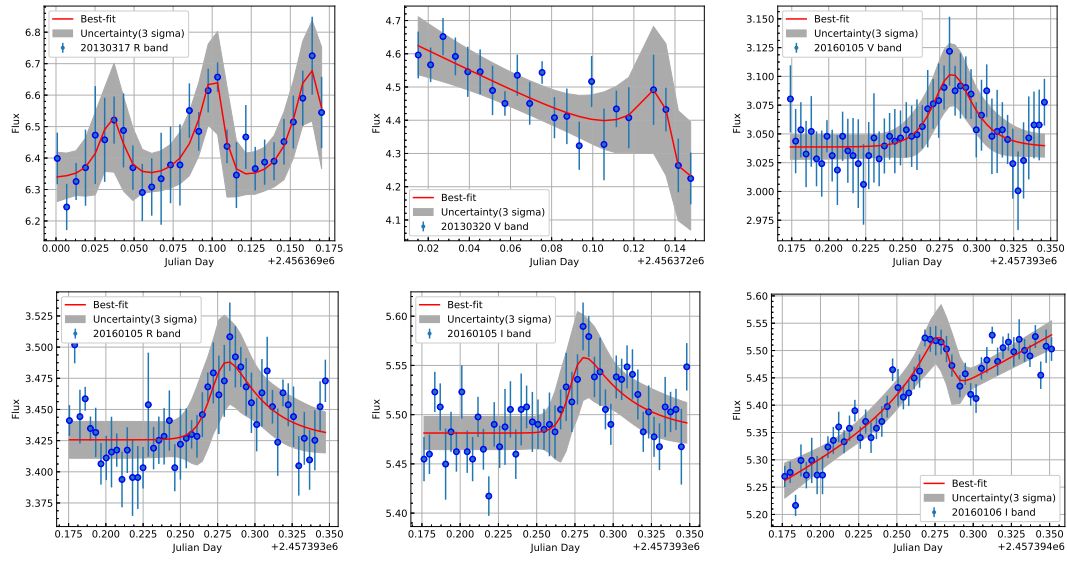
$$f_i(t) = \frac{2F_i}{\exp(\frac{t_i-t}{T_{r,i}}) + \exp(\frac{t-t_i}{T_{d,i}})} + F_0 \quad (7)$$

where  $F_0$  is the constant background flux level (could be a constant or linear model),  $F_i$  is the flare normalisation,  $t_i$  is the time of maximum flare,  $T_{r,i}$  is the flux rising timescale, and  $T_{d,i}$  is the flux decaying timescale.

Most of candidate light curves showed a monotonically increasing or decreasing trend, which meant the timescale of the variability was larger than the length of the data. Only two V light curves and four PV light curves showed complete flares. The results of the light curves are shown in Figure 5 and Table 4. On 17 March 2013, the durations of each flare were about 18.5, 21.1 and 23.1 min, respectively. The minimal variability timescale, defined as the minimum of the rising or decaying timescale, was about 5.3 min. On 20 March 2013, the duration was 24.3 min and the minimal variability timescale was about 2.7 min (with  $\pm 4.688$  min error), which is unreliable. On 5 January 2016, the durations were about 34.4, 45.6 and 51.3 min in the V-, R- and I-bands, respectively. The corresponding minimal variability timescales were about 14.6, 9.2 and 6.5 min. On 6 January 2016, the flux showed a monotonic increase in the V-, R- and I-bands, and in the I-band, there was a significant flare with the flux increasing. The duration was about 42.7 min and the minimal variability timescale was about 4.7 min. This was the identifiable minimum timescale in the whole observational campaign.

**Table 4.** The fitting parameters and their errors of Figure 5.

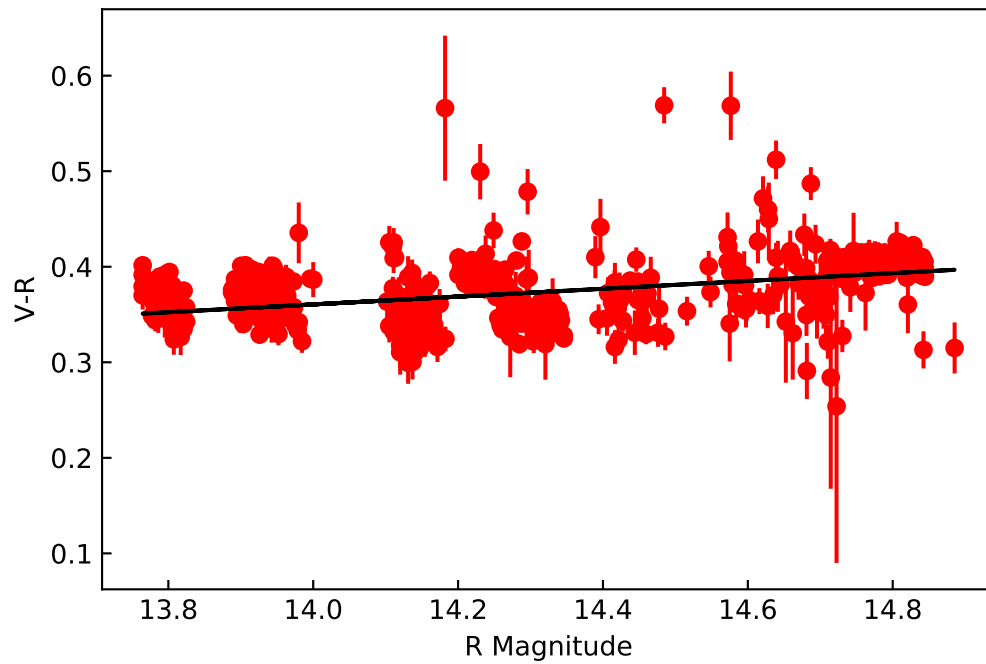
Date	Band	ID	$F_i$	$t_i$ +2,456,000	$T_{r,i}$ min	$T_{d,i}$ min
17 March 2013	R	1	$0.192 \pm 0.052$	369.037	$10.612 \pm 4.861$	$7.916 \pm 3.970$
		2	$0.291 \pm 0.048$	369.104	$15.788 \pm 3.865$	$5.326 \pm 1.983$
		3	$0.342 \pm 0.049$	369.164	$13.585 \pm 3.228$	$9.547 \pm 3.588$
20 March 2013	V	1	$0.159 \pm 0.061$	$372.136 \pm 0.002$	$21.615 \pm 12.064$	$2.712 \pm 4.668$
5 January 2016	V	1	$0.063 \pm 0.009$	$1393.282 \pm 0.007$	$14.594 \pm 6.451$	$19.763 \pm 7.906$
	R	1	$0.052 \pm 0.014$	$1393.275 \pm 0.006$	$9.160 \pm 6.090$	$36.480 \pm 16.328$
	I	1	$0.056 \pm 0.015$	$1393.274 \pm 0.004$	$6.536 \pm 4.596$	$44.752 \pm 17.941$
6 January 2016	I	1	$0.079 \pm 0.016$	$1394.282 \pm 0.003$	$37.986 \pm 14.515$	$4.672 \pm 2.890$



**Figure 5.** Light curves of OJ 287 on 17 March 2013 for R-band; 20 March 2013 for V-band; 5 January 2016 for V-band; 5 January 2016 for R-band; 5 January 2016 I-band; and 6 January 2016 for I-band. The lines are the best fitting by an asymmetric flare template.

### 3.2. Colour Index

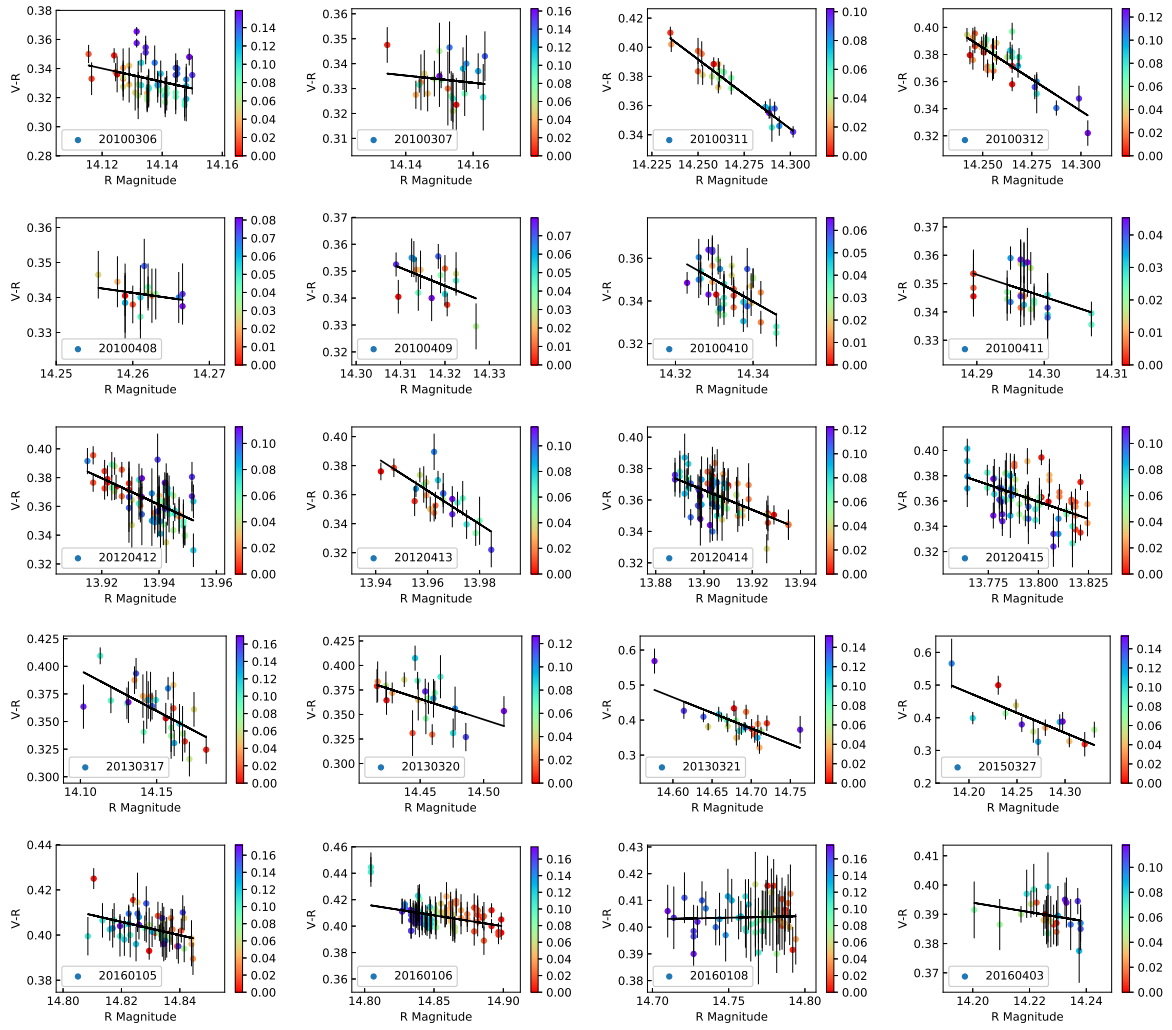
It is common that brightness variations are often associated with changes in spectral shapes. Thus we also investigated the correlations between the colour indices and magnitude variations. Here we consider whether variations in the  $V - R$  colour indices of OJ 287 changed with respect to variations in its brightness in the  $R$ -band on short- and long-term bases. These colour and magnitude plots of OJ 287 are displayed in Figures 6 and 7. We calculated the best linear fit as shown by the straight lines in Figures 6 and 7 for the colour index against magnitude, and the slope, intercept, linear Pearson correlation coefficients and corresponding null hypothesis probability values are also listed in Table 5. A positive Pearson's  $r$  coefficient between the colour index and apparent magnitude of the blazar implies a positive correlation, which means the source tends to be bluer when it is brighter (BWB) or redder when fainter (RWF), while a negative Pearson's  $r$  coefficient suggests the opposite correlation: redder when brighter (RWB) or bluer when fainter (BWF) behaviour. We found a weak positive correlation between the colour indices and the  $R$ -band magnitude on long timescales during these observations of OJ 287 (shown as Figure 6). Then we measured Pearson's  $r$  coefficient between the colour indices and the  $R$ -band magnitude on intra-night timescales, as plotted in Figure 7. Ten nights showed a significant negative correlation between the colour indices and  $R$ -band magnitude on intra-night timescales, with probability values below 0.001; other nights showed nearly achromatic (ACH) behaviour. On 11–12 March 2010 and 12–13 April 2012, the OJ 287 colour-magnitude correlations tended to show BWF behaviour, while on 14–15 April 2012 they tended to show RWB behaviour.



**Figure 6.** Optical colour index vs magnitude plots of OJ 287 during our monitoring. Fitted value for slope is  $0.041 \pm 0.003$ , intercept is  $-0.212 \pm 0.039$  and Pearson's  $r$  coefficient is 0.450 with  $\text{Pro } 6.080 \times 10^{-46}$ .

**Table 5.** Color-magnitude dependencies and colour-magnitude correlation coefficients on short timescales of Figure 7.

Date	Data Points	Slope	Intercept	Pearson's Coefficient	Pro
6 March 2010	49	$-0.460 \pm 0.188$	$6.833 \pm 2.661$	$-0.336$	0.018
7 March 2010	22	$-0.145 \pm 0.235$	$2.382 \pm 3.324$	$-0.137$	0.545
11 March 2010	22	$-0.954 \pm 0.063$	$13.989 \pm 0.895$	$-0.959$	0.000
12 March 2010	35	$-0.935 \pm 0.108$	$13.702 \pm 1.539$	$-0.833$	0.000
8 April 2010	16	$-0.311 \pm 0.295$	$4.774 \pm 4.210$	$-0.271$	0.310
9 April 2010	16	$-0.671 \pm 0.346$	$9.950 \pm 4.951$	$-0.460$	0.073
10 April 2010	40	$-1.026 \pm 0.245$	$15.047 \pm 3.506$	$-0.562$	0.000
11 April 2010	29	$-0.783 \pm 0.292$	$11.547 \pm 4.171$	$-0.459$	0.012
12 April 2012	69	$-0.918 \pm 0.147$	$13.157 \pm 2.053$	$-0.606$	0.000
13 April 2012	24	$-1.151 \pm 0.209$	$16.434 \pm 2.917$	$-0.762$	0.000
14 April 2012	75	$-0.623 \pm 0.106$	$9.033 \pm 1.473$	$-0.567$	0.000
15 April 2012	71	$-0.541 \pm 0.104$	$7.827 \pm 1.440$	$-0.529$	0.000
17 March 2013	27	$-0.730 \pm 0.179$	$10.689 \pm 2.528$	$-0.633$	0.000
20 March 2013	21	$-0.419 \pm 0.182$	$6.421 \pm 2.626$	$-0.468$	0.033
21 March 2013	24	$-0.891 \pm 0.179$	$13.471 \pm 2.628$	$-0.728$	0.000
27 March 2015	14	$-1.225 \pm 0.293$	$17.866 \pm 4.182$	$-0.770$	0.001
5 January 2016	49	$-0.296 \pm 0.095$	$4.795 \pm 1.411$	$-0.413$	0.003
6 January 2016	50	$-0.157 \pm 0.056$	$2.742 \pm 0.839$	$-0.373$	0.008
8 January 2016	51	$0.011 \pm 0.041$	$0.234 \pm 0.600$	$0.040$	0.778
3 April 2016	25	$-0.161 \pm 0.117$	$2.679 \pm 1.659$	$-0.276$	0.181



**Figure 7.** Optical colour index vs magnitude plots of OJ 287. The colour bars indicate the progression of time in units of  $JD_{start}$ . Detail about the colour-magnitude correlation variability can be found in Table 5.

#### 4. Discussion and Conclusions

Blazars, being dominated by a relativistic jet oriented at a small angle from the line of sight, display variability ranging from a few hundredths of a magnitude to more than several magnitudes. Their variability timescales range from minutes to years. Either factors intrinsic to the jet or due to changes in overall jet power could explain such variability. The shock-in-jet model can be used to explain most of variabilities in blazars [10,27–30]. Any changes in the physical mechanisms of blazars, such as to the magnetic field, electron density or velocity, can trigger a shock that leads to flares when propagating along the relativistic jet. Turbulence behind the shock-in-jet can be a good way to explain the fast variability in the IDV of blazars [10,31–33]. However, when blazars are in a low state, the emission from the accretion disk dominates the jet emission. Instability in the accretion disk, such as at “hot spots”, can also produce IDVs [9,10,34–36]. Perturbations on the accretion disks transfer into the relativistic jet and, being Doppler amplified, might also produce fluctuations [9,10,37,38].

Previous colour index studies suggest apparent dichotomy responses: BL Lacs become BWB and FSRQs become RWB [9,10,39–44]. The BWB behavior in BL Lacs can be interpreted as electrons being accelerated to preferentially higher energies before radiatively cooling, and the RWB behavior in FSRQs is explained by the addition of redder, non-thermal jet emission to an already bluer, thermal disk

component. These results are often preferentially observed during flaring, which causes the selection effect that the two extreme ends of the colour-magnitude correlation are most regularly reported. Ruan et al. [45] point out that spectral variability could be the result of hot spots in the accretion disk emission by using a sample of 604 variable quasars. Isler et al. [46] suggest colour-magnitude variability should be a continuum rather than a dichotomy. They use empirical data from 3C 279 to explain how blazars can smoothly evolve from a jet-quietest, disk-dominated colour profile to an actively jet-dominated state and associated colour profile and back to a jet-quietest state. Our results of colour-magnitude correlations on long and short timescales tend to be interpreted in the framework of a relatively bluer accretion disk and redder jet change in intensity. In addition, the jet intensity could change in a short timescales, while the disk intensity changed more slowly. Thus the intra-night colour behavior was alternatively BWF, RWB or ACH, which can be described by the variation in contributions of the non-thermal relativistic jet in OJ 287.

We carried out multi-band optical photometric monitoring of the blazar OJ 287 over 34 observation nights between 6 March 2010 and 3 April 2016. We searched for flux and colour-magnitude variation on IDV timescales. The IDV amplitudes of OJ 287 were small for the observed flux variabilities. The blazar was in a highly active and variable state during January 2016, showing relatively large amplitude flux variations:  $F_{var}$  of 1.3–2.1%, and Amp of 5.8–9.0%. OJ 287 showed a weak positive correlation on the long timescale, while negative correlations were found on IDV timescales. Our present analysis cannot distinguish between the possible physical mechanisms in detail. Very dense and highly precise simultaneous multi-band observations are necessary. Thus OJ 287 should continue to be monitored whenever possible.

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**Conflicts of Interest:** The authors declare no conflict of interest.

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